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apparent distance has nearly its maximum value for a few days on either side of the dates given.

I have recently examined *Capella* again on two good nights. At the time of the last examination, August 2, 1901, the seeing was excellent and all the powers mentioned above were successively used. Glass color-screens of different shades were used at times to reduce the light. At the time of this examination the zenith distance of *Capella* varied from 45° to 38° . With all powers the central disc of the star's image was sharply defined, very small, and apparently perfectly round.

W. J. HUSSEY.

TIMES OF THE ELONGATIONS OF *MIMAS*.

While making measurements of the relative positions of the inner satellites of *Saturn* during the present summer, I noticed that *Mimas* reached its greatest eastern and western elongations about 4.2 hours earlier than the predicted times given in the American Ephemeris. On writing to Dr. HARSHMAN, Director of the Nautical Almanac Office, in reference to the matter, he informed me that the predicted times were based upon Professor HALL's elements. He also furnished me the systematic corrections necessary to reduce the times of elongation as derived from Professor HALL's elements to those derived from Dr. H. STRUVE's elements, as follows:—

| | |
|--------|----------------------|
| 1900.0 | — 3 ^b .68 |
| 1901.0 | — 4 .02 |
| 1902.0 | — 4 .37 |
| 1903.0 | — 4 .72 |
| 1904.0 | — 5 .06 |

The times of elongation of *Mimas* for the years 1900, 1901, 1902, and 1903, as given in the American Ephemeris, are based upon HALL's elements. Those for 1904 are based upon STRUVE's elements. From my measures it appears that STRUVE's elements represent the motion closely.

W. J. HUSSEY.

DUPLICITY OF 66 *TAURI* AND OF THE PRINCIPAL COMPONENT OF Σ 2339.

While examining 66 *Tauri* with the 12-inch telescope, February 24, 1900, I discovered that it is a close pair. Several months

passed before I measured it, and it is only recently that second and third measures have been obtained. These measures give —

| | | | | |
|----------|--------|-------|-----|-----|
| 1900.920 | 22°.7 | 0".26 | 5.8 | 6.0 |
| 1901.584 | 205 .3 | 0 .24 | 6.0 | 6.0 |
| .658 | 23 .6 | 0 .25 | 6.0 | 6.0 |

The magnitudes of the two components are so nearly the same that I have not been able to determine the quadrant with certainty.

66 Tauri is a naked-eye star. Its proper motion has been determined, and, according to AUWERS, is given by the equations —

$$\mu = - 0''.0027, \quad \mu' = - 0''.004.$$

On account of the character of the pair and by reason of its proper motion, it is probable that *66 Tauri* is a binary.

I have recently found the principal component of Σ 2339 to be a close pair. The star did not appear to be quite round to me with the 12-inch telescope. On this account I examined it with the 36-inch refractor, which clearly showed it double. I have made the following measures of it:—

| | | | | |
|----------|-------|-------|-----|-----|
| 1901.603 | 85°.2 | 0".20 | 8.0 | 8.2 |
| .608 | 87 .0 | 0 .18 | . . | . . |
| .623 | 86 .7 | 0 .20 | 8.0 | 8.2 |

The components of this pair are a little unequal in brightness, so that the quadrant is readily determinable.

I have referred these stars to Professor BURNHAM, who has kindly looked them up for me, and states that they are both new.

W. J. HUSSEY.

OBSERVATIONS OF THE SPECTROSCOPIC BINARY η PEGASI.*

The binary character of η Pegasi was discovered in August, 1898, from observations made with the Mills spectroscope, and announced in the *Astrophysical Journal* for October, 1898. It was the first one of some thirty-five spectroscopic binary systems discovered in the past three years with this efficient instrument. As a basis for determining the orbit of the bright component of

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